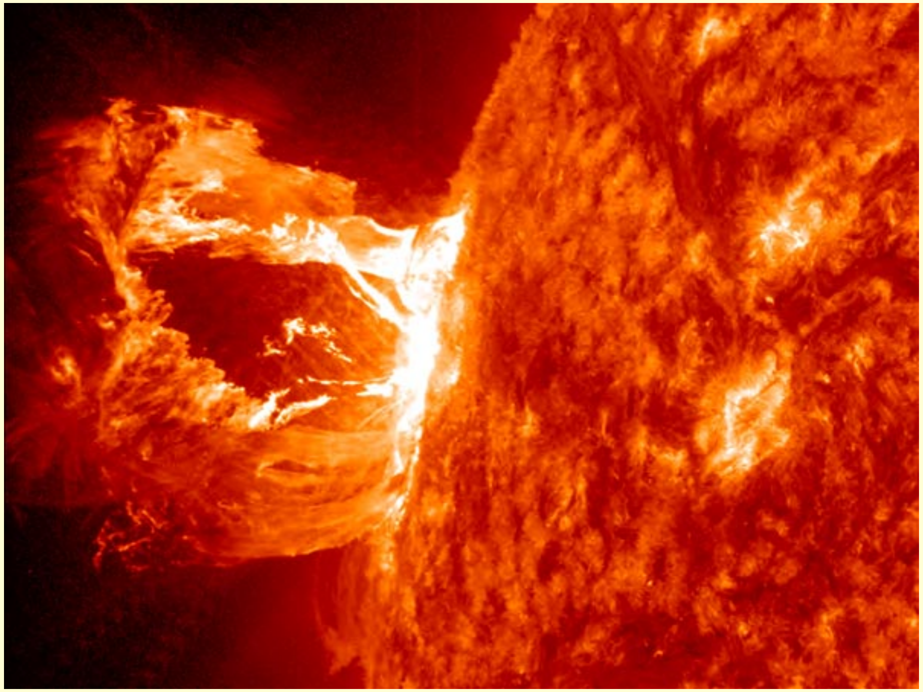
Solar Magnetism



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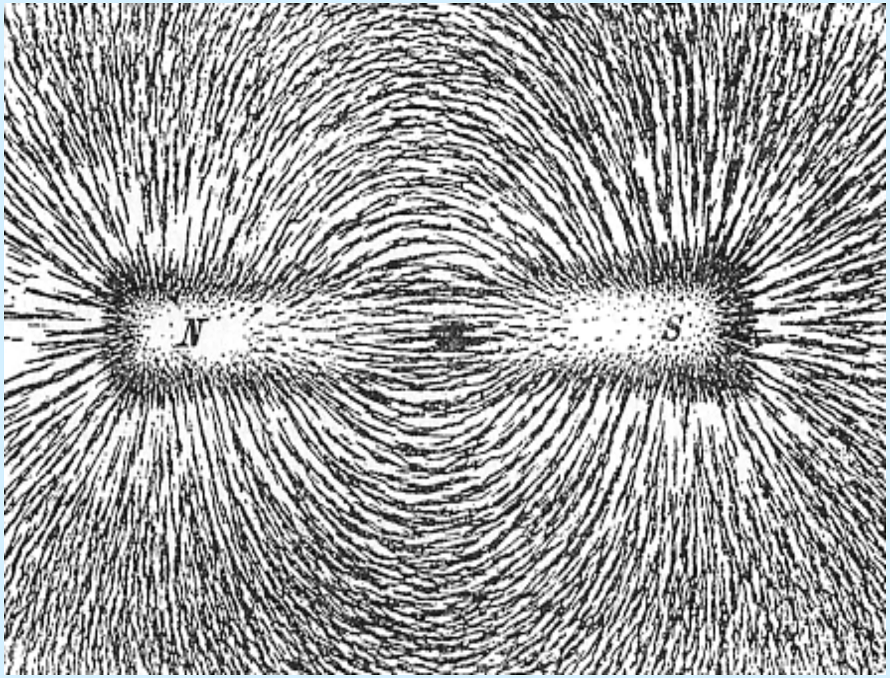
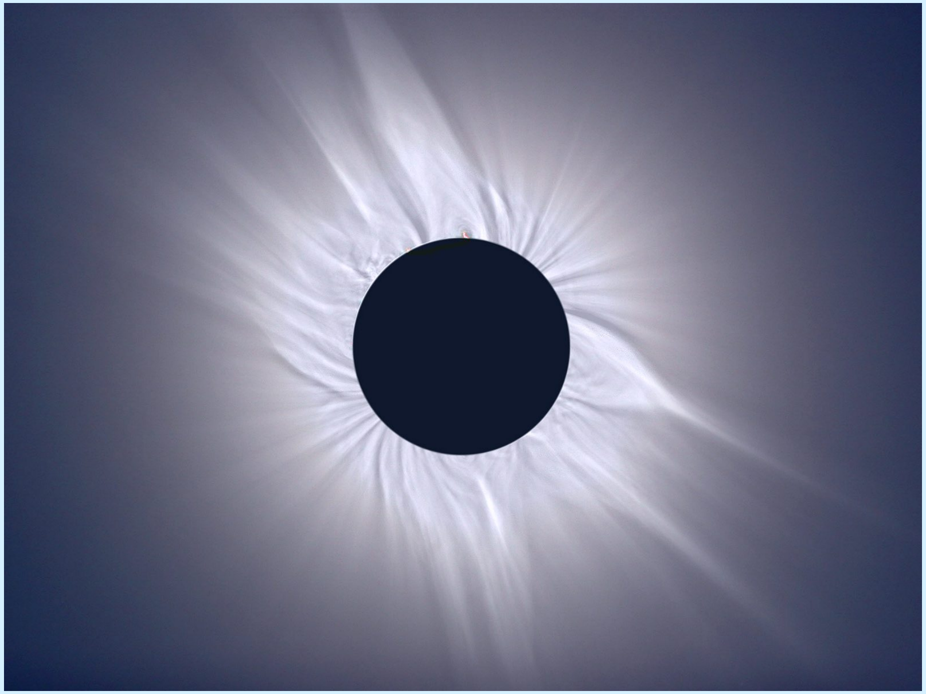


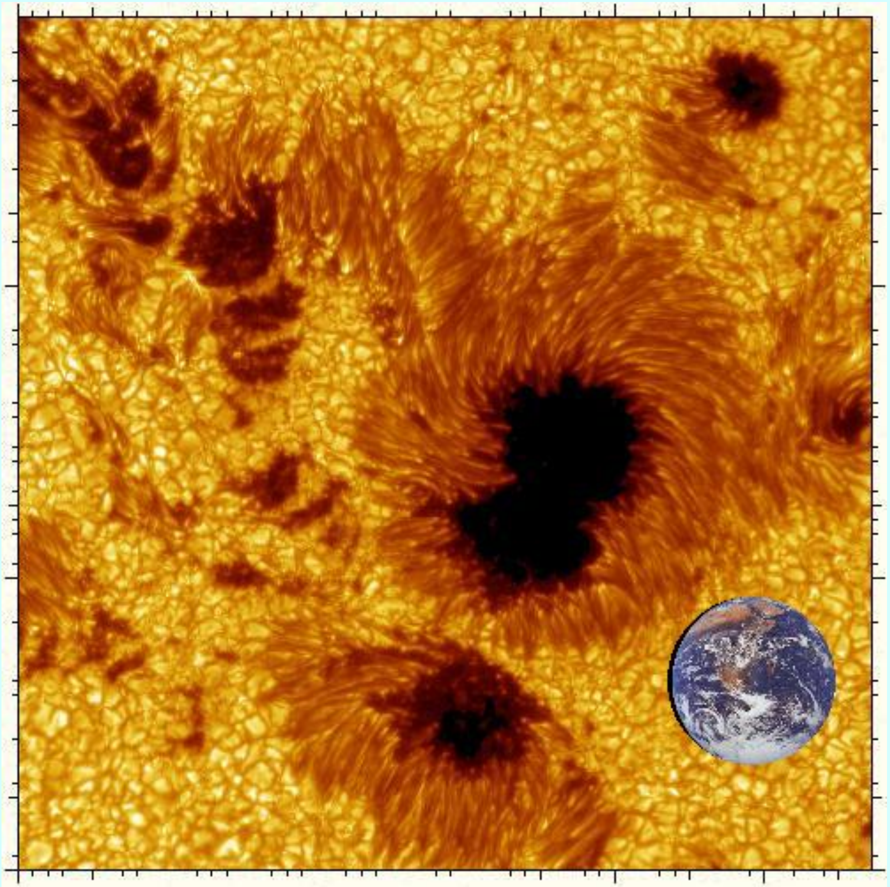


Iron filings around a bar magnet

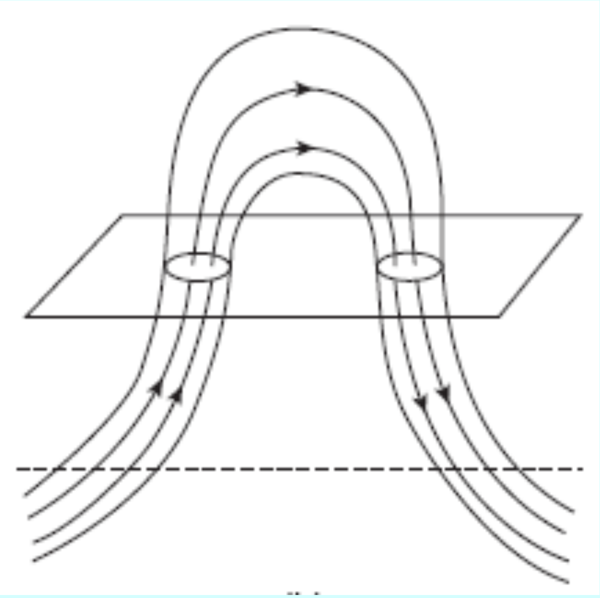
Solar corona during a total solar eclipse

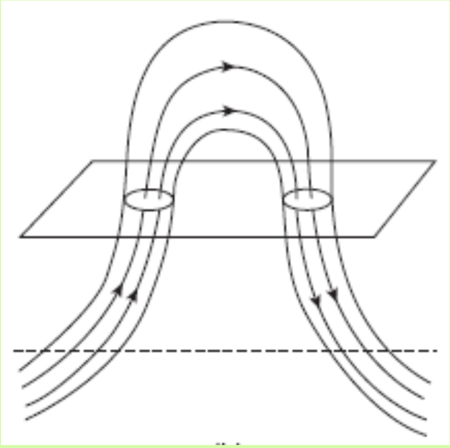
Solar magnetic fields do affect our lives!

Hale (1908) discovered magnetic fields in sunspots from Zeeman splitting of magnetic fields (0.3 tesla)
Hale et al. (1919) – Often two large sunspots are seen side by side with opposite polarities



A strand of magnetic flux has come through the surface!

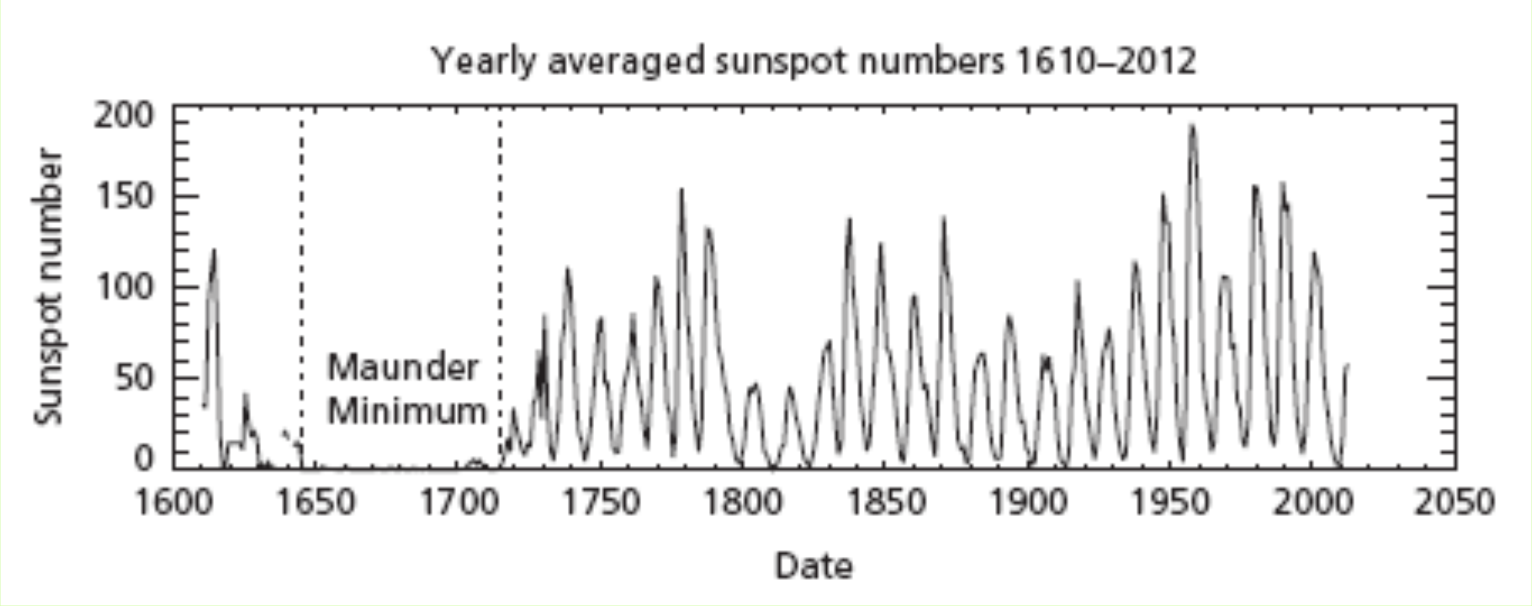




The solar corona is full of such magnetic loops

Loops emit in X-ray and Extreme ultra-violet (EUV)

1844: Heinrich Schwabe discovered the 11-year sunspot cycle

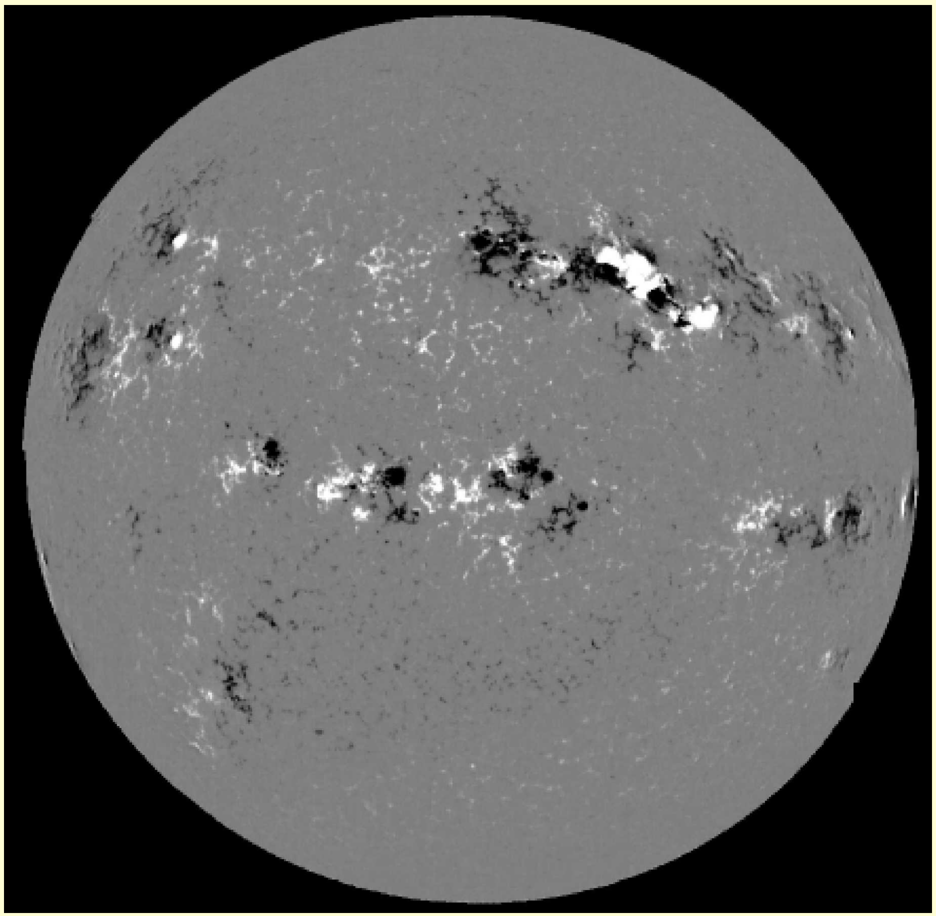


The cycle has many irregularities

Maunder minimum 1640 – 1720

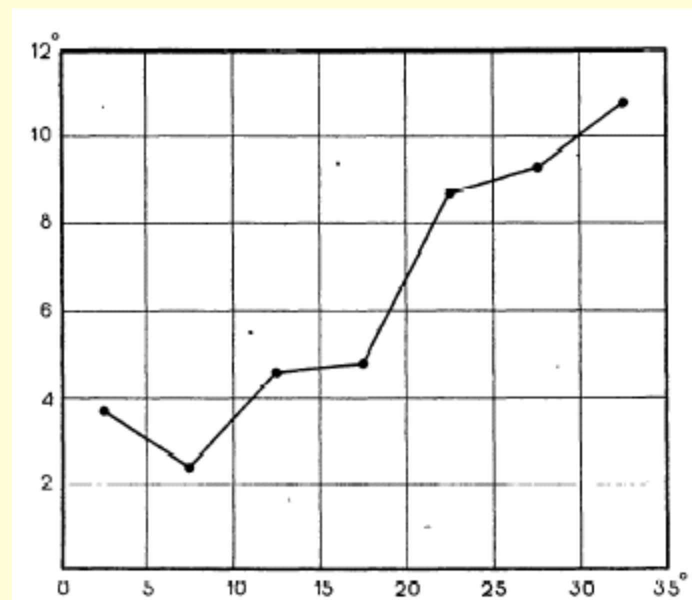
Little ice age!

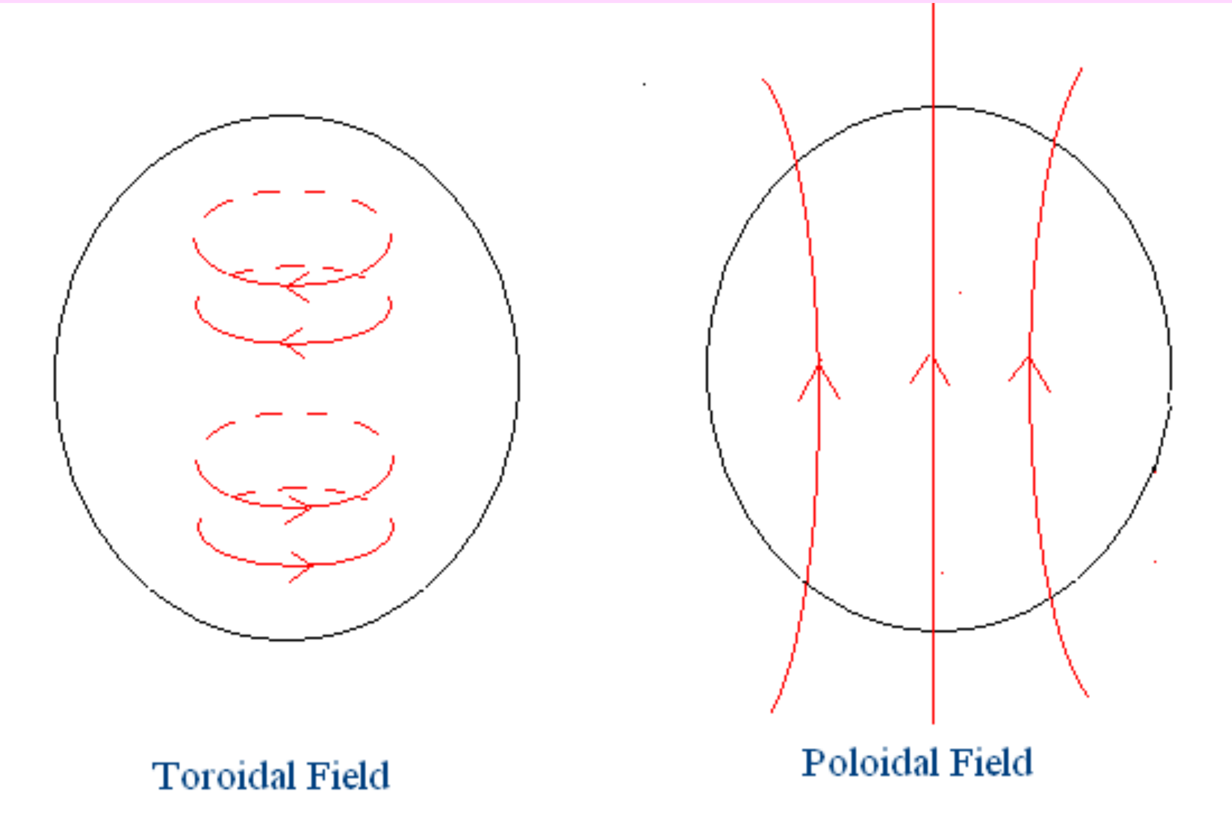
Magnetogram map (white +ve, black –ve)
Polarity is opposite (i) between hemispheres; (ii) from one 11-yr cycle to next >> 22-yr period



Tilt of bipolar regions increases with latitude

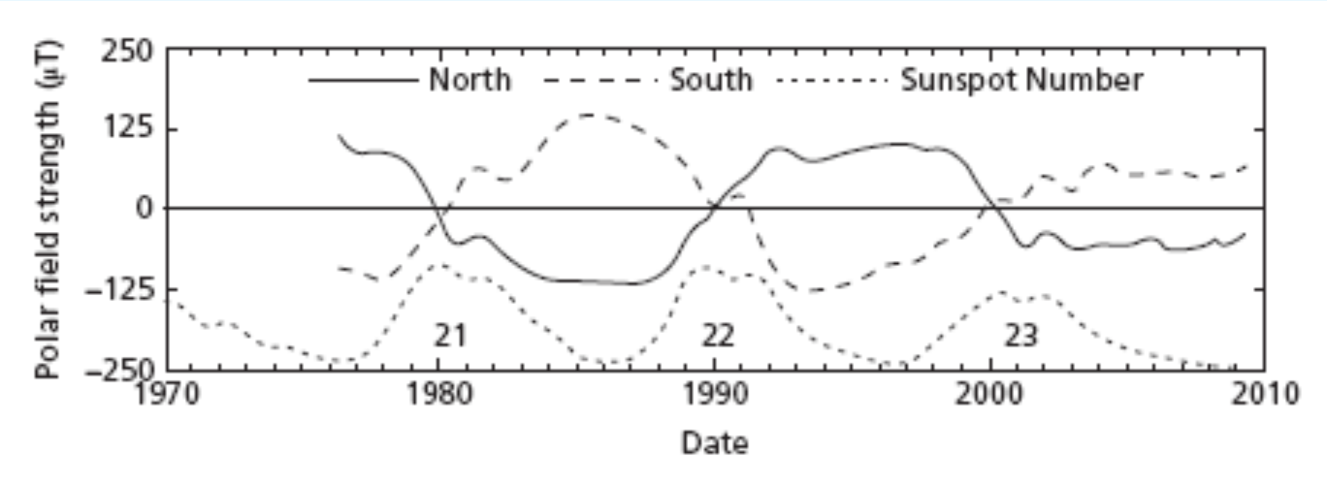
- Joy's law (Joy 1919)





Parker (1955) suggested oscillation between the toroidal and poloidal fields.

Babcock & Babcock (1955) detected the weak poloidal field (~ 10^{-3} T)



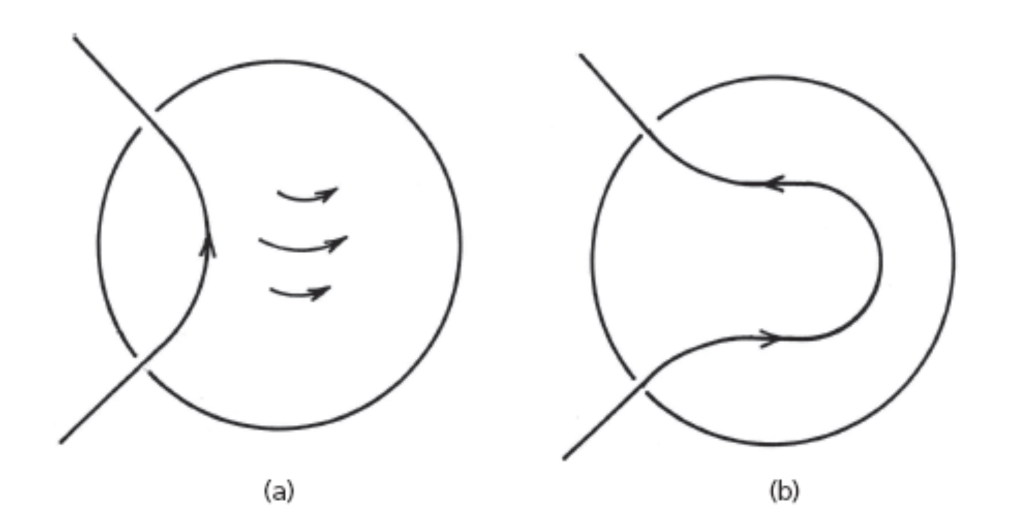
The polar fields and the sunspot number as functions of time

There is indeed an oscillation between toroidal magnetic field (indicated by sunspot number) and poloidal magnetic field (indicated by polar field), as envisaged by Parker (1955)

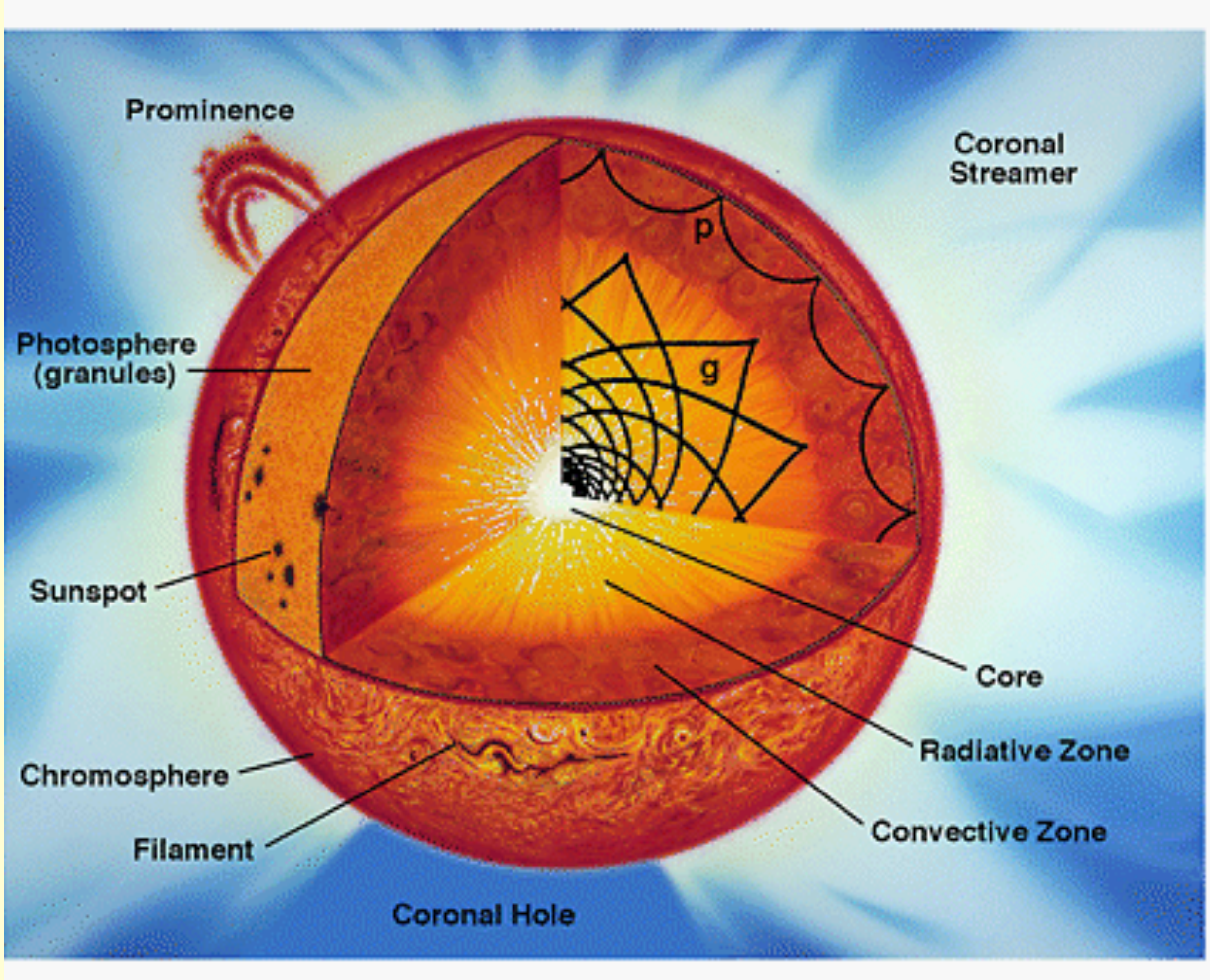
Central Dogma of solar dynamo theory

- The Sun's magnetic field consists of both a toroidal component and a poloidal component
- Sunspots arise out of the toroidal field
- The polar field is a manifestation of the poloidal field
- There is a process producing toroidal field from poloidal field
- There is a process producing poloidal field from toroidal field, so that we get the cycle

The Sun rotates around its axis in about 27 days, but does not rotate like a solid body
Equatorial regions rotate faster => Differential rotation



The toroidal magnetic field is stretched by the differential rotation to generate the poloidal magnetic field



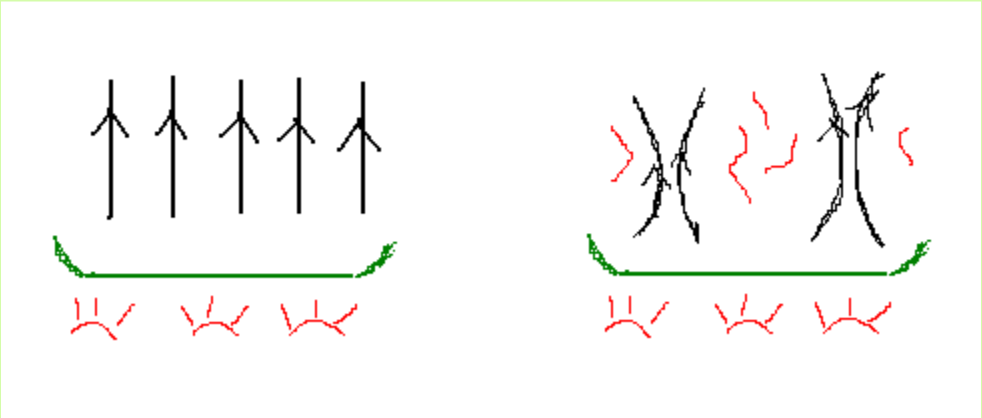
Sunspots are magnetic field concentrations in turbulent plasma

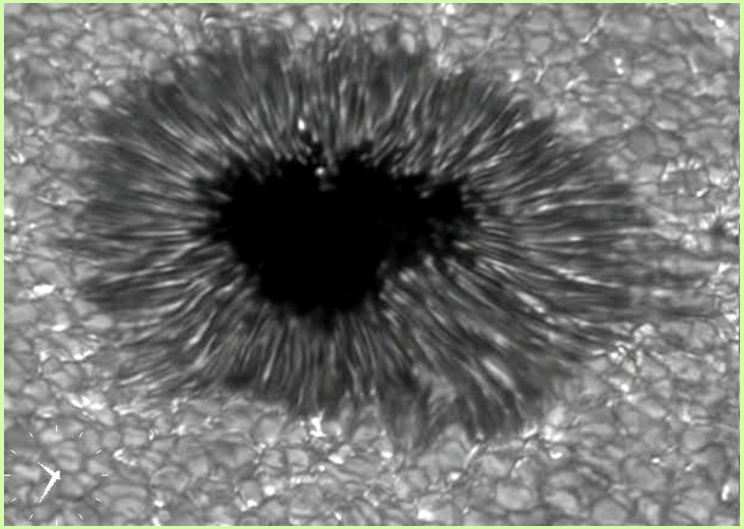
Heat is transported outward by radiative transfer till $0.7R$ and then by convection in the zone $0.7R – R$

Magnetoconvection

Linear theory – Chandrasekhar 1952

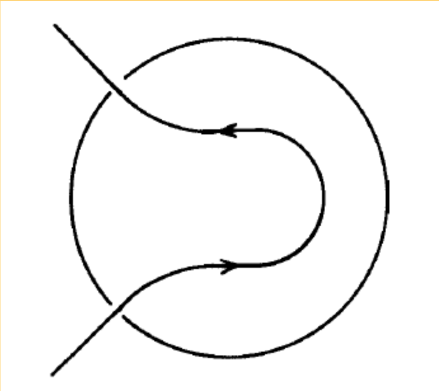
Nonlinear evolution – Weiss 1981; . . .





Sunspots are magnetic field concentrations with suppressed convection

Magnetic field probably exists as flux tubes within the solar convection zone



Why do parts of the toroidal magnetic field float up?

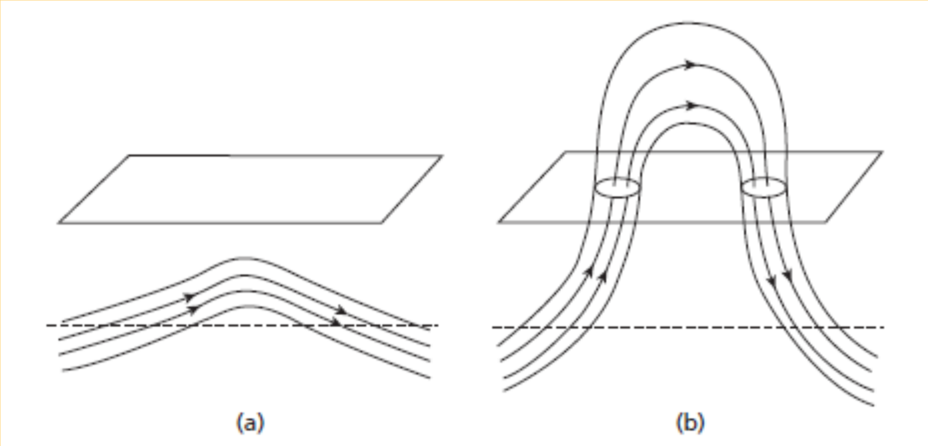
Horizontal Magnetic Flux Tube

$$p_{\rm out} = p_{\rm in} + \frac{B^2}{2\mu} \longrightarrow p_{\rm in} \leq p_{\rm out}$$

Usually the inside is under-dense

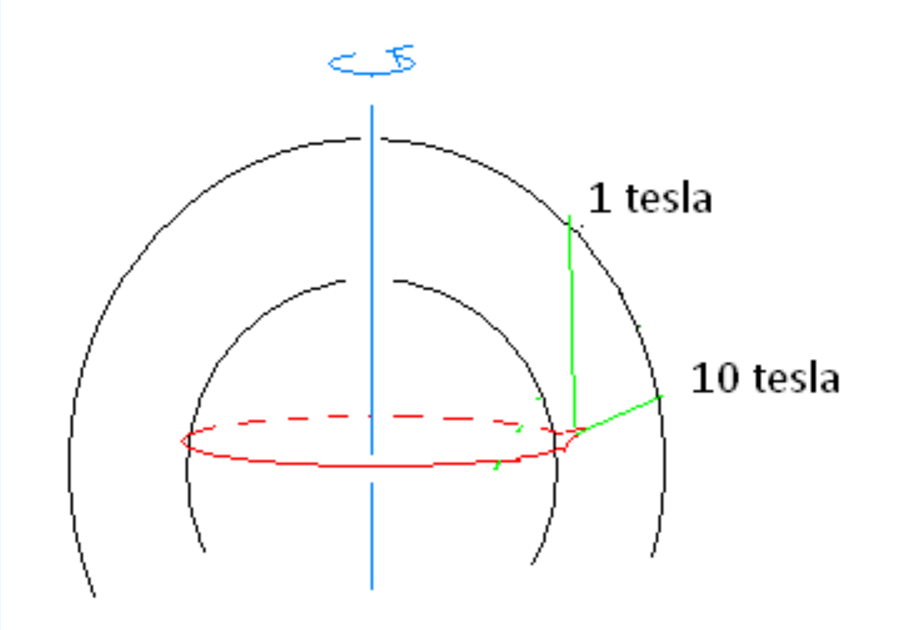
Magnetic buoyancy (Parker 1955)

Very destabilizing within the convection zone, but much suppressed below its bottom



3D dynamics of flux tubes in solar convection zone

(Choudhuri & Gilman 1987; Choudhuri 1989; D'Silva & Choudhuri 1993; Fan et al. 1993; Caligari et al. 1995)



Early dynamo models suggested B at bottom to be 1 tesla, but such fields are diverted by Coriolis force (Choudhuri & Gilman 1987)

Only 10 tesla fields can emerge at sunspot latitudes

Early models of solar dynamo could not work with magnetic fields stronger than 1 tesla

Invoking some early ideas of Babcock (1961) and Leighton (1964), flux transport dynamo model for the sunspot cycle was developed to allow for much stronger fields

The workability of this model was demonstrated by Choudhuri, Schussler & Dikpati (1995)

Basic Equations

Magnetic field

$$\boldsymbol{B} = B(r,\theta)\boldsymbol{e}_\phi + \nabla \times [A(r,\theta)\boldsymbol{e}_\phi],$$

Velocity field

$$\Omega(r,\theta)\, r \sin\theta\, \mathrm{e}_\phi + \mathrm{v}$$

$$\frac{\partial A}{\partial t} + \frac{1}{s}(\boldsymbol{v}.\nabla)(sA) = \eta_{\mathrm{p}}\left(\nabla^2 - \frac{1}{s^2}\right)A + \alpha B,$$

$$\frac{\partial B}{\partial t} + \frac{1}{r}\left[\frac{\partial}{\partial r}(r v_r B) + \frac{\partial}{\partial \theta}(v_\theta B)\right] = \eta_{\mathrm{t}}\left(\nabla^2 - \frac{1}{s^2}\right)B + s(\boldsymbol{B}_{\mathrm{p}}.\nabla)\Omega + \frac{1}{r}\frac{\mathrm{d}\eta_{\mathrm{t}}}{\mathrm{d}r}\frac{\partial}{\partial r}(rB)$$

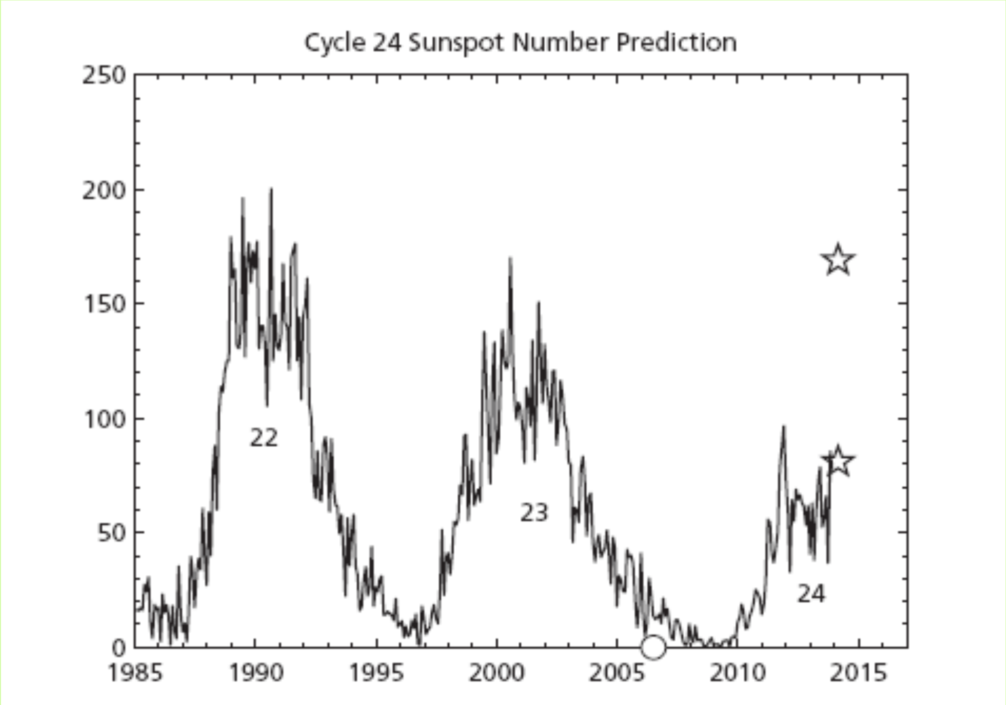
The code *Surya* solves these equations

For a range of parameters, the code relaxes to periodic solutions (Nandy & Choudhuri 2002, *Science* **296**, 1671)

Can we predict the strength of a sunspot cycle before its advent?

Dikpati & Gilman (2006) predicted a strong cycle 24

Choudhuri, Chatterjee & Jiang (2007) predicted a weak cycle 24

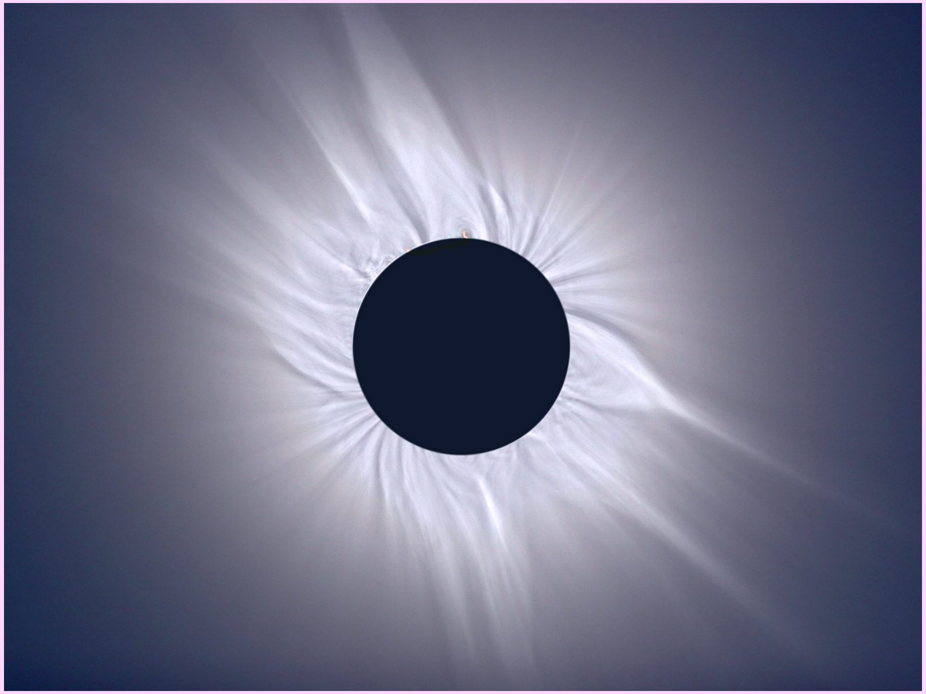


Our prediction is the first successful prediction of a sunspot cycle from a theoretical dynamo model!!

Although Sun's surface has a temperature of about 6000 K, the temperature of the corona is millions of degrees!!!

First inferred from spectral lines of the corona (Edlen 1943)

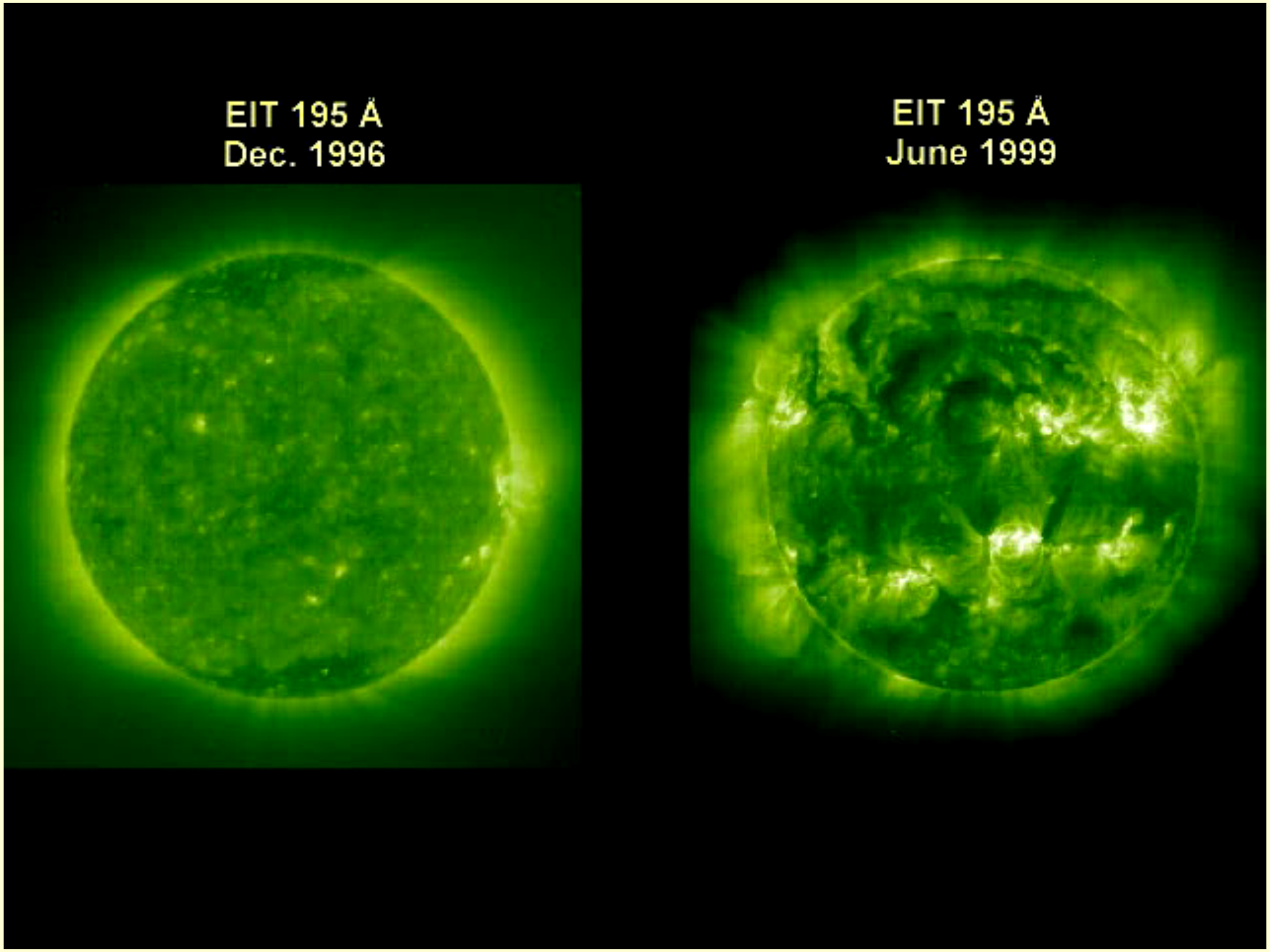
Do we have a violation of the second law of thermodynamics?



Hottest regions of the corona should emit X-rays and Extreme UV

Has to be detected from space

Rotating Sun seen in Extreme UV



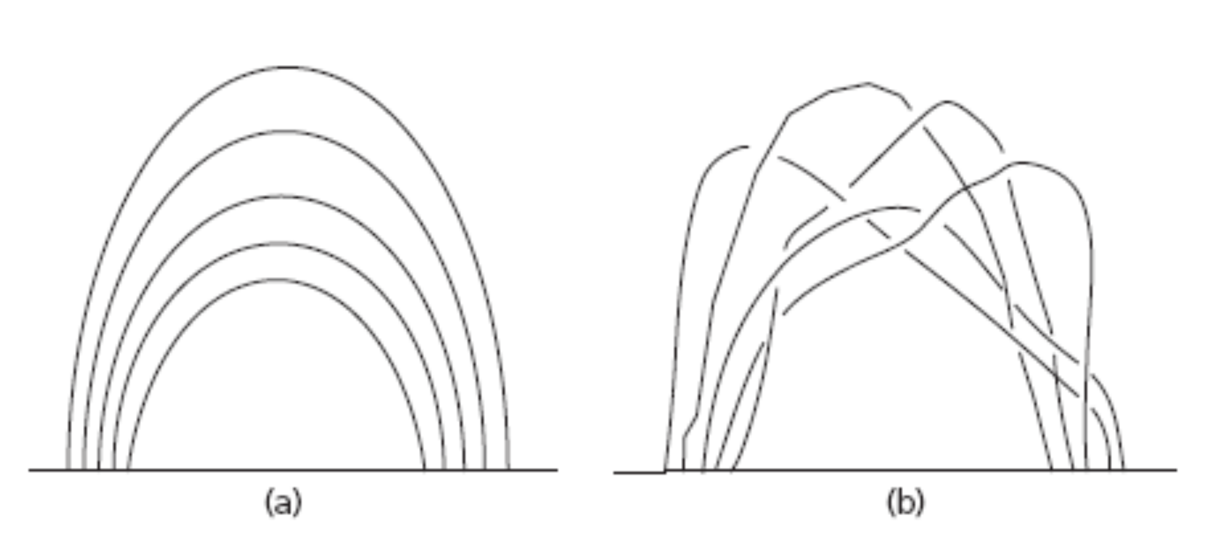
Solar Minimum

Solar Maximum

Magnetic fields in the coronal loops imply currents

Can the heat generated by the currents produce the high tempature of the corona? – Only if the currents flow through narrow regions!

Parker (1972) – disturbances of footpoints by convection tangles up magnetic fields

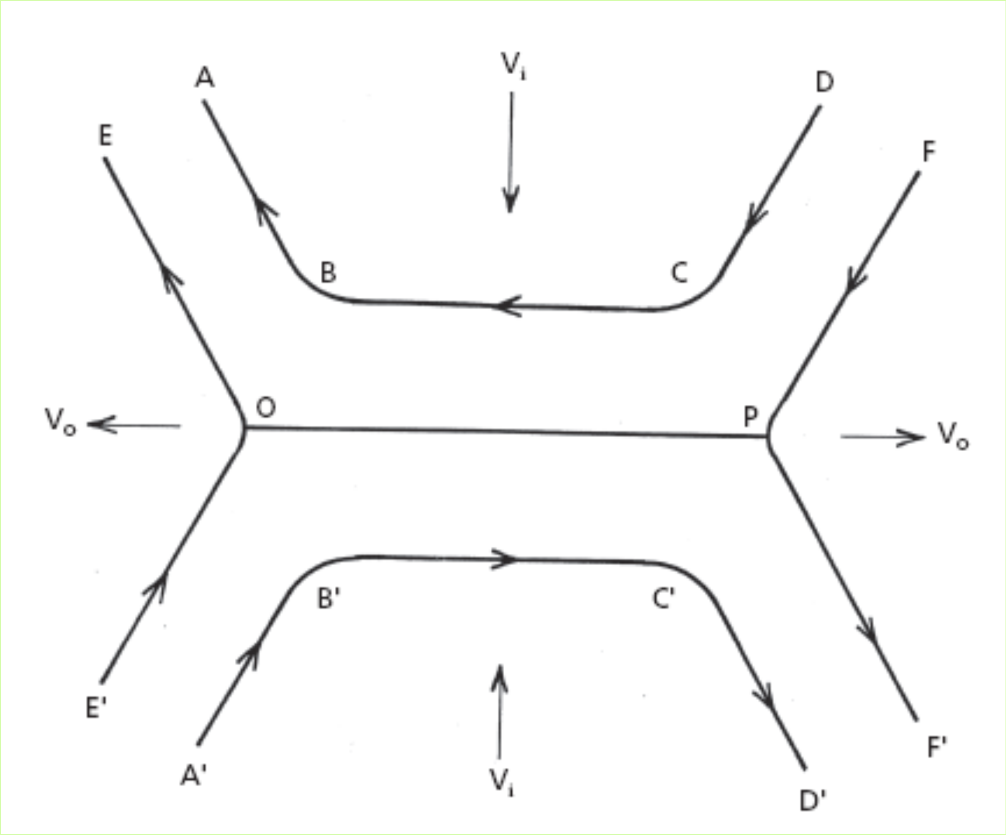


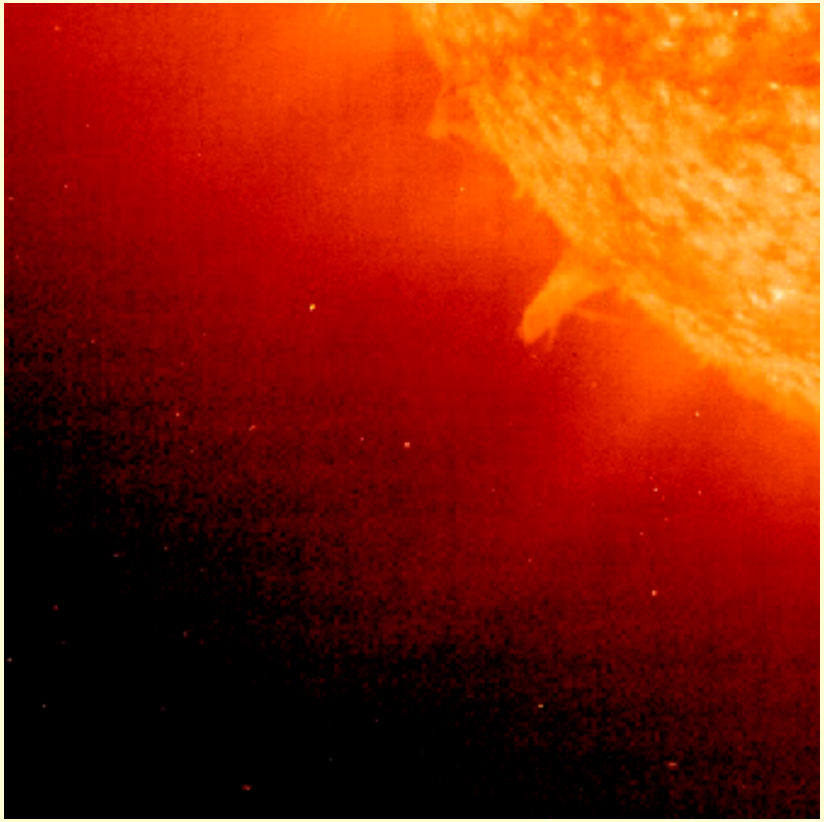
Currents flow through narrow regions and heat the corona

Magnetic reconnection – Sweet 1958, Parker 1957, Petschek 1964

Heating goes as j^2 and is important in the central region where current density j is high

Solar flares are caused by magnetic reconnection in gigantic regions where often two flux systems come across each other





Eruptive Prominence

Coronal Mass Ejection (CME)

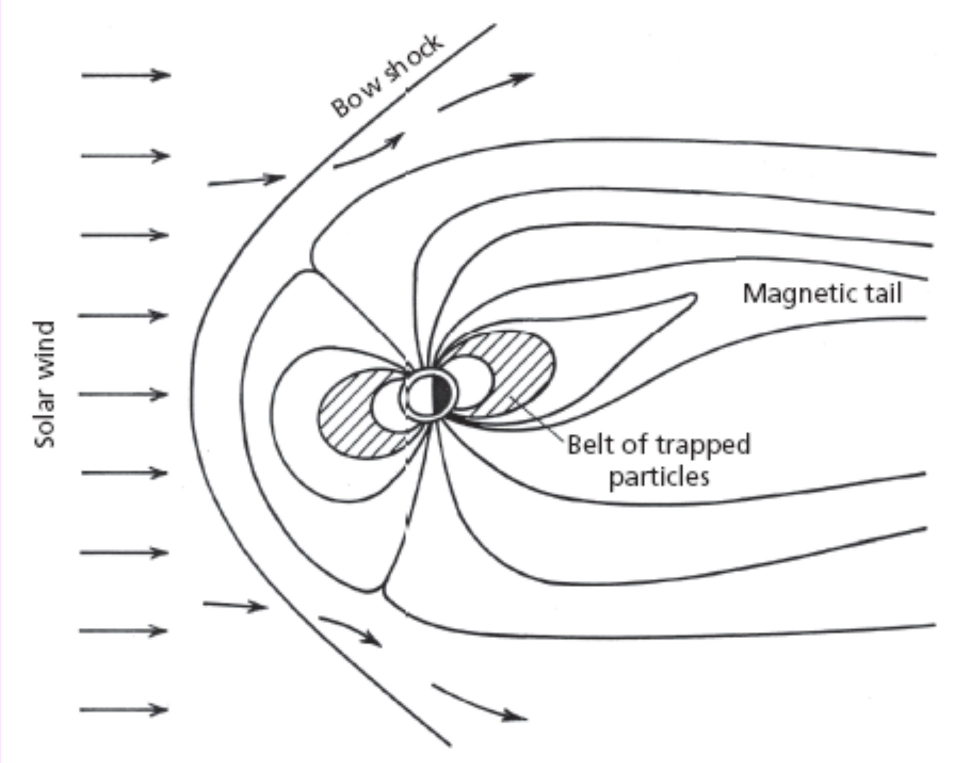
More flares, eruptive prominences and CMEs occur when there are more sunspots

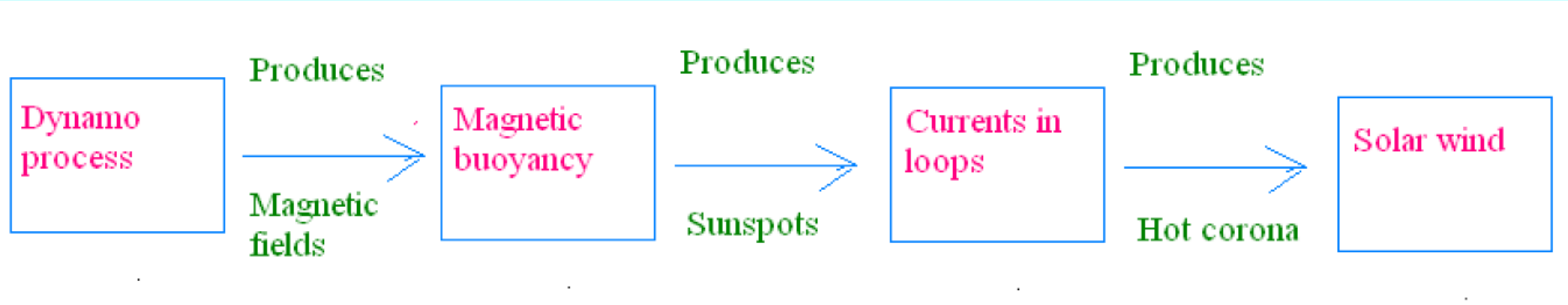
Parker 1958 – Sun's gravitational field is unable to keep the hot corona confined => a plasma outflow, the **solar wind**

Discovered by space missions within 3 – 4 years!

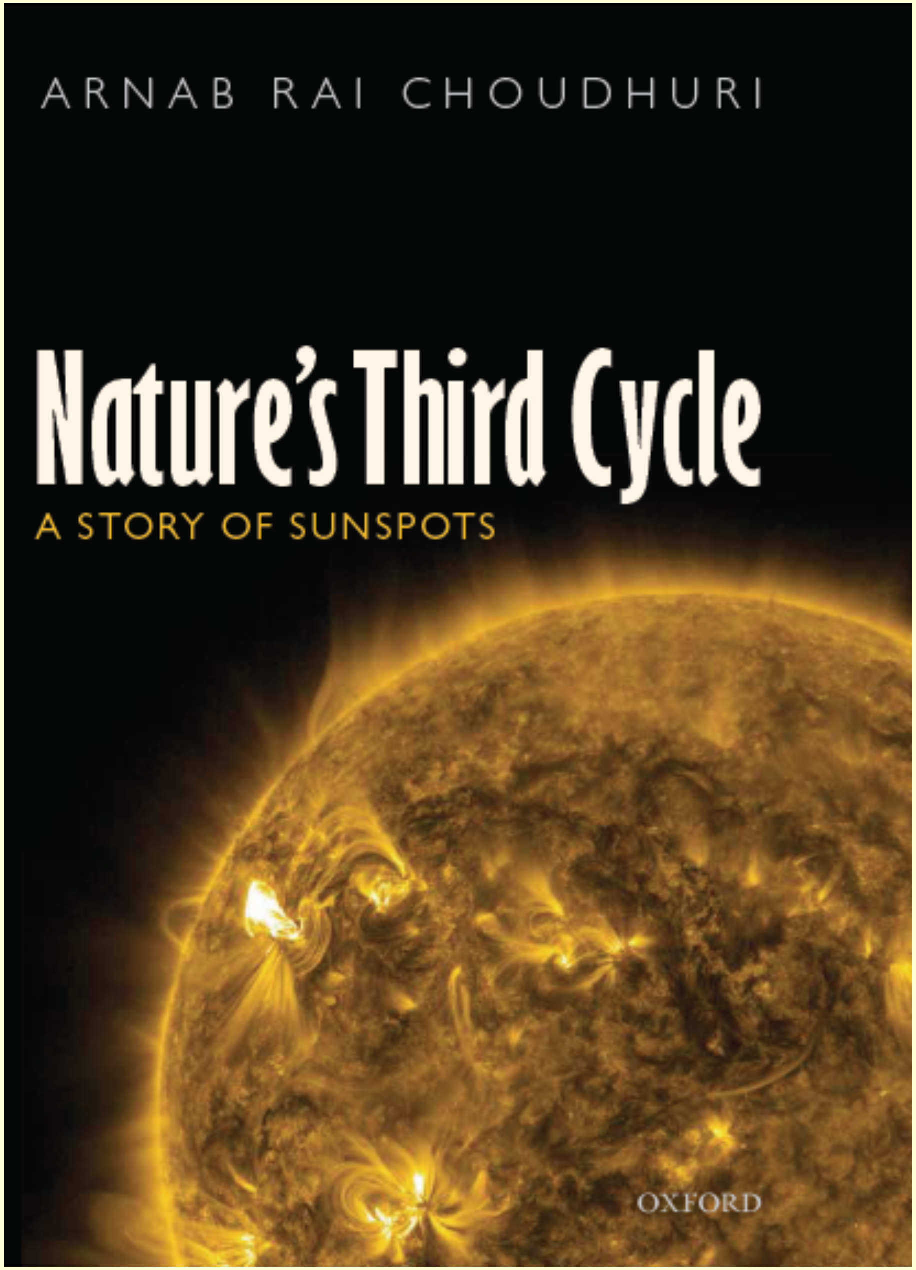
Solar disturbances may be carried with the solar wind, taking 3 – 4 days to reach the earth

The solar wind impinges on the Earth's magnetosphere, making it lop-sided





Eugene Newman Parker (1927 -)
My guru



My forthcoming popular science book to be published by Oxford University Press in January 2015

Amazon India is already selling Kindle edition for less than Rs. 300/-

Helioseimology

Leighton, Noyes & Simon 1962 – discovered solar oscillations

Deubner 1974 – recognized them as normal modes

Angular velocity distribution in the solar interior could be found by analyzing these oscillations

Strong differential rotation at the bottom of convection zone